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VERY LARGE TELESCOPE

GRAVITY Pipeline User Manual

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1 Introduction

This manual is a complete description of the data reduction recipes implemented for the GRAVITY pipeline, reflecting the status of the GRAVITY pipeline as of now.

The main part of the document is focused on the main feature of the pipeline useful to the science user of GRAVITY. The detailled appendixes may be more useful to ESO staff for the purpose of long term re-calibration and data quality control.

1.1 Reference and applicable documents

[1] *OIFits Explorer*. *http://www.jmmc.fr/oifitsexplorer*. 37

[2] *QFitsView*. *http://www.mpe.mpg.de/ ott/QFitsView/*. 37

[3] ESO/SDD/DFS, *http://www.eso.org/cpl/*. *CPL home page*. 54

[4] T. A. Pauls, J. S. Young, W. D. Cotton, and J. D. Monnier. A Data Exchange Standard for Optical (Visible/IR) Interferometry. *PASP*, 117:1255–1262, November 2005. 16, 38, 57

[5] C.Sabet P.Ballester. *VLTI Data Interface Control Document*. ESO, 1.0 edition, 3 June 2002. VLT-SPE-ESO-15000-2764. 37

2 GRAVITY Instrument Description

The GRAVITY instrument has been developed under ESO contract by the GRAVITY consortium.

The main purpose of the instrument is to measure the angular separation between stars in the vicinity of the Galactic Center black hole, and even the position and motion of the infrared emission of the black hole itself during flares.

GRAVITY recombines the light beams collected by either the four 8.2 m Unit Telescopes or the four VLTI Auxiliary Telescopes. The general principle of the instrument is to recombine the light coming from one (single field mode) or two (dual field mode) astronomical targets in two different beam combiners: the Fringe Tracker (FT) and the Science Combiner (SC). In single field mode, the light from the target is split between the FT and SC channels using a beam splitter, while in dual field mode a mirror directs the light from each target to their respective beam combiners. The FT is optimized to record fringes at very high frequency (up to 1 kHz), in order to measure and compensate in real time the atmospheric piston effect using a dedicated actuator in the instrument. As the observed targets are both within the atmospheric isoplanetic patch, the correction of the atmospheric piston by the FT stabilizes the fringes of the SC channel. This gives the possibility to integrate for up to several tens of seconds, and therefore reach a high sensitivity, even at relatively high spectral dispersion, on the SC channel.

The properties of the interference fringes are measured separately in the FT and SC beam combiners. GRAVITY measures the classical interferometric observables of any source, as the previous VLTI instruments (VINCI, MIDI, AMBER and PIONIER). The FT spectral resolution is limited to 5 spectral channels over the K band. The SC has three available spectral resolutions: low (40), medium (400) and high (4000), providing approximately 10, 200 and 1800 spectral channels over the K band.

The phases of the SC beam combiner are referenced to the FT using a metrology system that encompasses the optics of the VLTI up to the secondary mirror of the telescopes. Thanks to this link between the two beam combiners, GRAVITY provides very accurate measurements of the differential position of the fringe pattern obtained, for each baseline, between a reference star (in the FT channel) and the target star (in the SC channel). Given the VLTI FOV, the angular separation between the reference object and the science target is limited to 5 arcsec with the ATs and 2 arcsec with the UTs. Within this restricted separation, the final accuracy on the relative astrometry is expected to be of a few tens of μarcsec.

3 Quick start

3.1 The concept of recipe

The pipeline is built over "recipes" that implement the basic steps of the reduction. The recipe input is a Set Of Files (SOF), which contains the list of files to reduce, associated with the required calibrations. Each recipe has a list of possible options. Each recipe produces one or several products. The recipes can be triggered by several mechanisms:

gasgano is the ESO user interface to classify, select files (= construct a SOF) and launch recipes.

esorex is the ESO command line tool to launch a recipe on any SOF manually written by the user.

reflex is the ESO environment allowing to reduce an entire directory in batch mode.

The GRAVITY consortium has also built a dedicated python script which browses the content of the current directory, classifies the files, writes the SOFs and automatically triggers the corresponding recipes via esorex.

3.2 The concept of SOF

A SOF is a consistent list of files to be sent to a recipe. Each file in the list shall be assigned a *DO Category*, which tells the recipe its purpose in the reduction. In the GRAVITY pipeline, the DO categories are trivially built from the DPR.TYPE, DPR.CATG and PRO.CATG keywords in the FITS header.

3.3 Instrument calibrations

In order to reduce interferometric observations, it is mandatory to calibrate the detectors and the combiners. These calibrations are obtained via dedicated observations of the internal source with all shutters closed, one shutter open at a time, two shutters open at a time, and all shutters open.

The recipe **gravity_dark** creates the DARK calibration product, which contains the mean detector bias and the detector readout noise. It shall match the detector and the optical setup of the observation.

The recipe **gravity_p2vm** creates the BAD (bad pixel), FLAT (internal transmission), P2VM (internal phase and contrast) and WAVE (wavelength map) calibration products. They shall match the optical setup of the observation for the SC and the FT, as well as the detector gain for the FT.

These products are all needed in order to reduce the interferometric observations of science target.

3.4 From raw data to raw visibilities

The first step is to reduce the raw interferometric observations into uncalibrated measurement of the visibilities and closure phases. This step is generally done file-per-file, that is each OBJECT exposure in the raw directory-has a corresponding files in product directory (although it is possible to reduce several files together). Note that this step includes the averaging over all the frames that compose a standard exposure. This step is rather long (several hours for an entire night).

The recipe is **gravity_vis**.

The product (*_VIS) is an uncalibrated OIFITS file, containing the flux, the square visibilities, the complex visibilities and the closure phases for both the SC and the FT combiners.

The recipe produces an optional product, called P2VMRED, which contains many intermediate signals of the processing. It is very useful to assess the data quality. It is also possible to restart the reduction from the P2VMRED product, thus saving time for users who want to explore several values of recipe parameters.

3.5 From raw visibilities to calibrated visibilities

The second step is to calibrate the observation of science objects with the observation of calibration stars. This step could be performed “globally”, that is all the uncalibrated *_VIS oifits can be loaded first, then the pipeline searches for consistent sequences (same setup, same DIT, same wavelength table...), and then each sequence is calibrated independently with the calibration stars that could be found within it. This step is rather short (1min for an entire night).

The recipe is **gravity_viscal**.

The product (*_VIS_CALIBRATED) is a calibrated OIFITS file, ready for science.

4 Instrument Data Description

4.1 RAW science data

The RAW frames created when observing have the following DPR.TYPE:

OBJECT,DUAL are observations of a nearby pair of objects, one feeding the fringe-tracker (FT) and the other feeding the science combiner (SC). It can be of category SCI or CAL.

OBJECT,SINGLE are observations of a single object, feeding both the fringe-tracker (FT) and the science combiner (SC). It can be of category SCI or CAL.

SKY,SINGLE
SKY,DUAL are observation of an empty patch of the sky near the object in order to measure the sky brightness.

The OBJECT,DUAL and OBJECT,SINGLE types have a category DPR.CATG=SCI when observing a science target, and a category DPR.CATG=CAL when observing a calibration star used to monitor the transfer function.

4.2 RAW calibration data

The RAW frames used to calibrate the instrument on a daily-basis have the following DPR.TYPE:

DARK are observations with all shutters closed, in order to calibrate the detector dark level and the detector + dark level noise.

FLAT are observations of the internal source with one shutter open, in order to calibrate the positions of the spectra on the detectors and the internal transmission of the instrument.

P2VM are observations of the internal source with two shutters open, in order to calibrate the internal contrasts and phases of the instrument.

WAVE
WAVESC are observations of the internal source with all shutters open, in order to calibrate the wavelength table, and the internal closure phases. The WAVE data are recorded by scanning on both SC and FT, with FDDL in open loop and no fringe tracking (in order to calibrate the FT wavelengths). The WAVESC data are recorded by scanning only SC, and FDDL in close-loop and fringe tracking (to calibrate the SC wavelengths).

4.3 STATIC calibration

The STATIC calibration frames have the following DPR.TYPE:

DISP_MODEL is the model of the optical dispersive index $n(\lambda)$ of the fiber differential delay lines (FDDL) of the instrument.

DISP_VIS is an intermediate product when building DISP_MODEL, used to visualise the quality of the FDDL stretching sequence.

DIAMETER_CAT is the catalog of stellar diameters used to estimate the transfer function.

EOP_PARAM is a list of Earth Orientation Parameters (EOP) and DUT1 versus time. These corrections are only needed for the most demanding astrometric measurements.

4.4 PRODUCT calibration data

The PRODUCT of the calibration by the recipes **gravity_dark** and **gravity_p2vm** are identified by the following PRO.CATG keyword:

DARK contains images with the dark level and variance for the SC and FT detectors.

BAD contains images with the identified bad pixels for the SC and the FT detectors.

FLAT contains images of the profiles used to extract the SC spectra from the detector. There is one extracted spectrum per output of the detector and per polarisation if split (thus 24 or 48 spectra for each SC and FT combiners).

WAVE contains tables with the effective wavelengths of each channel of every spectra extracted with the profile. These tables are necessary to re-align the different spectra (outputs of the detector) onto a common wavelength grid.

P2VM contains tables with the internal transmission, contrast and phase of every output of the detector versus wavelength. These form the so-called pixel-2-visibility matrix used to extract the interferometric visibility from the spectra.

4.5 PRODUCT science data

The products of the science reduction by the recipe **gravity_vis** are identified by the following PRO.CATG keywords:

SINGLE_SKY **DUAL_SKY**	contains the mean brightness of the sky and its variance. Contrarily to DARKs, these quantities possibly depend on the instrument mode because different optics are used.
SINGLE_SCI_VIS **SINGLE_CAL_VIS** **DUAL_SCI_VIS** **DUAL_CAL_VIS**	are OIFITS files [4] with the uncalibrated flux, squared visibilities, complex visibilities and closure phases extracted from the raw observation of an object. SCI/CAL corresponds to a science target or a calibration star used to monitor the transfer function.
SINGLE_SCI_P2VMRED **SINGLE_CAL_P2VMRED** **DUAL_SCI_P2VMRED** **DUAL_CAL_P2VMRED**	are the raw data already processed through the P2VM algorithm. They contain flux per beam, and coherent flux per baseline, for each individual frame of the exposure. As such, they are intermediate products between the RAW data and the final, averaged, OIFITS. They also contain many intermediate results of the processing. The file size is huge (>200Mb). It is meant to assess the overal data quality and tune the reduction parameters. It is not used for science. Its format is inspired by OIFITS, but it is not strictly compliant.
SPECTRUM **PREPROC**	contains the RAW data already corrected for cosmetic and collapsed into one spectrum per combiner output. In SPECTRUM, the data are not yet re-interpolated into a common spectral wavelength grid, while this step is done in PREPROC. As such, they are debug-level intermediate products between the RAW data and the final, averaged, OIFITS.

The PRODUCT of the final calibration step by the recipe **gravity_viscal** are identified by the following PRO.CATG keywords:

SINGLE_SCI_VIS_CALIBRATED **DUAL_SCI_VIS_CALIBRATED**	are the final OIFITS file of the reduction, science ready. They contain the interferometric observations calibrated with the transfer function.
SINGLE_CAL_TF **DUAL_CAL_TF**	are OIFITS files containing the transfer function value estimated by the corresponding observation of a calibration star. It is the observed visibility of a calibrator divided by its visibility estimated from its diameter.
SINGLE_SCI_TF **DUAL_SCI_TF**	are OIFITS files containing the estimated (interpolated) value of the transfer function at the time of the corresponding science target observation.

5 Data Reduction

5.1 Graphical overview of the cascade

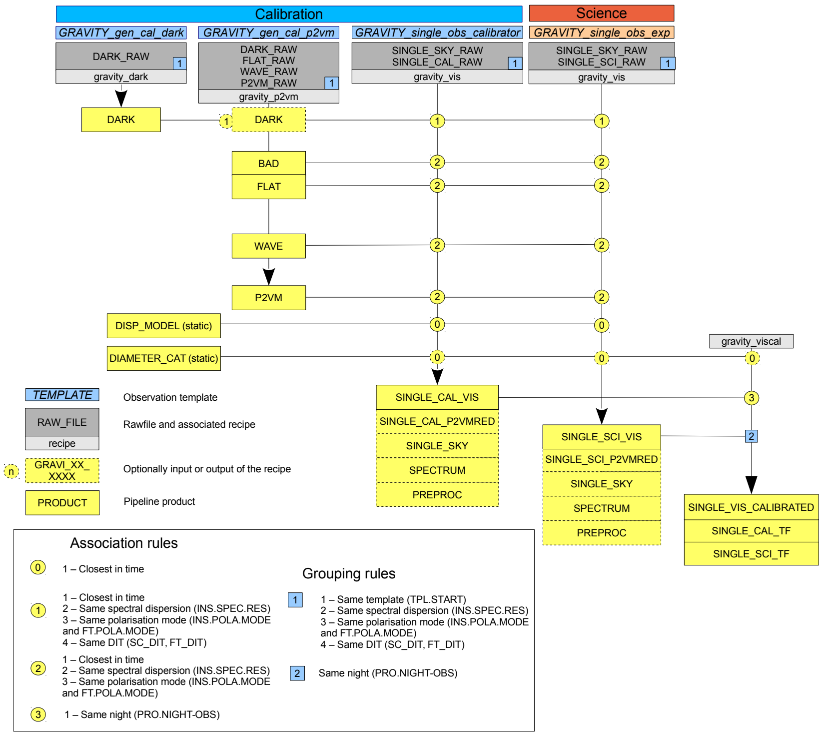

Figure 5.1: Data reduction cascade for observations in mode SINGLE. As of now, the cascade is the same for the mode DUAL.

5.2 Using Gasgano

Gasgano, provides a graphic interface for data browsing, classification and association, and offers several other utilities such as easy access to recipes documentation and preferred data display tools.

Gasgano can be started from the system prompt in the following way:

```
gasgano &
```

Use the *Add/Remove Files* entry of the *File* menu to load data. The data are hierarchically organised as preferred by the user. More information about a single frame can be obtained by clicking on its name.

Frames can be selected from the main window for being processed by the appropriate recipe. Before launching the recipe, its configuration may be modified on the *Parameters* panel (on top). The window contents might be saved for later use by selecting the *Save Current Settings* entry from the *File* menu.

Please refer to the *Gasgano User's Manual* [7] for a more complete description of the *Gasgano* interface.

5.3 Using EsoRex

EsoRex is a command line utility for running pipeline recipes. It may be embedded by users into data reduction scripts for the automation of processing tasks. Users are free to define manually the input SOF and the appropriate configuration parameters.

A SOF for *EsoRex* is a simple ASCII file listing the files and their *DO Category*. Examples of SOF are given in the description of each recipe. Note that '#' is the comment character.

The basic syntax to use *EsoRex* is the following:

```
esorex [esorex_options] recipe_name [recipe_options] set_of_frames.sof
```

To get more information on how to customise *EsoRex* (see also [7]), or on a specific recipe, run the commands:

```
esorex  -h
esorex  -h  recipe_name
esorex  --man-page recipe_name
```

For more information on *EsoRex*, see `http://www.eso.org/cpl/esorex.html`.

5.4 Using run_gravi_reduce.py python script

The consortium has written a simple python script that classifies the files in the current directory, associates them with calibrations, and runs the corresponding recipes.

Instrument calibration and data reduction is triggered by the following script:

```
run_gravi_reduce.py [options]
```

The script performs the following steps:

1. Trigger the recipe **gravity_dark** on all standalone DARK.
2. Trigger the recipe **gravity_p2vm** on all sequences of 1 x DARK, 4 x FLAT, 6 x P2VM, 1 x WAVE, 1 x WAVESC.
3. Trigger the recipe **gravity_vis** on all OBJECT,* files.

The script will search for (static) calibrations in the current directory if any. It will also search in the upper directory `../common_calibration/`. The products are created in the `reduced/` subdirectory.

Before triggering a recipe, the script writes the corresponding SOF and the esorex command in the `reduced/` subdirectory. Thus one can manually reproduce a given reduction by executing:

```
./reduced/GRAVI.XXXX_esorex.sh
```

To get more information on syntax and options, run the commands:

```
run_gravi_reduce.py -h
```

Transfer function and trending can be then triggered with the following script:

```
cd reduced/
run_gravi_trend.py [options]
```

The script performs the following steps:

1. Trigger the recipe **gravity_viscal** on all *CAL_VIS.
2. Trigger the recipe **gravity_viscal** on all *SCI_VIS.
3. Produce trending plots to check the transfer function.

The products are created in the `calibrated/` and `trend/` subdirectories.

6 Known Problems

6.1 Spectral calibration

The spectral calibration between baselines is accurate to 0.1nm. It corresponds to half a pixel in spectral direction (thus 1/4 of the spectral resolution element) in HR.

This uncertainty can generate biases in the closure phase, which amount to $\approx$3deg when observing at a group-delay of 40 μm.

The absolute spectral calibration is accurate to 0.5nm, which corresponds to one spectral resolution elements in HR.

6.2 Uncertainties in products

The uncertainty of product data contains the statistical noise only, computed by bootstrapping over the NDIT samples when possible. It does not contain the calibration uncertainty.

When the number of valid DIT within an exposure is lower than 5, the statistic to compute the final error bars also include additional MonteCarlo realisation of the *theoretical* photon and detector noise (to reach 5 samples). These uncertainties are thus less realistic.

7 Pipeline Recipe Interfaces

7.1 List of all recipes

We here list the role of each recipe. The input, output, options and QC parameters are detailled in each dedicated subsection.

gravity_badpix	Detect the bad pixels on the detectors.
gravity_biasmask	*Not Offered* Determine which pixels can be used to measure the bias of SC detector.
gravity_dark	Calibrate the detector noise and background level.
gravity_disp	Calibrate the linearity and the dispersion of the differential delay lines.
gravity_eop	Download the last values of the Earth Orientation Parameters and DUT from IERS.
gravity_image	*Not Offered* Reconstruct an image from visibilities.
gravity_nab	*Not Offered* Calibrate the narrow angle baseline.
gravity_p2vm	Detect the bad pixels on the detectors, calibrate the wavelength tables, calibrate the interferometric contrast and phase.
gravity_piezo	*Not Offered* Calibrate the response of the piezo actuators.
gravity_postprocess	Post-process the products, to fine-tune their content.
gravity_vis	Compute the visibilities from raw observation of OBJECT.
gravity_vis_from_p2vmred	Compute the visibilities from P2VMRED intermediate product.
gravity_viscal	Calibrate visibilities with the transfer function (atmospheric interferometric response).
gravity_wavelamp	Measure the position of the Argon lines in the spectra.

7.2 gravity_dark

This recipe computes the DARK calibration for the SC, the FT and the ACQ detectors. The SC detector is first debiased using the biaspixels, before computing the dark mean and rms. For detectors, the mean dark level of each pixel and the stdev of each pixel are saved in the output product.

1. Loop on input dark files and concatenate them
2. Compute the median and rms of these concatednated files
3. Save the product (FT, SC, ACQ camera into same product)

Input

DO.CATG	
DARK_RAW	raw dark, all shutters closed (DPR.TYPE=DARK)

Output

PRO.CATG	
DARK	dark calibration

Parameters

Name	short description
–static-name	Use static names for the products (for ESO). [FALSE]
–bias-method	Methode to average the biaspixels when cleaning-up the SC detector (only apply to MED and LOW). Ideally the same value shall be used when reducing the DARK with gravity_dark and the OBJECT with gravity_vis. <MEDIAN | MEDIAN_PER_COLUMN> [MEDIAN]

Quality control

QC in DARK	short description
PIXBIAS AVG	Mean of the pixels used to removed the detector bias. This value shall be added to MEDIANDARK SC to trend the detector dark illumination [adu].
PIXBIAS RMS	Standard deviation over the pixels used to remove the detector bias [adu].
MEDIANDARK ACQ	Median of the dark level in the acquisition camera detector [adu]
MEDIANDARK SC	Median of the dark level in the Science Combiner detector [adu]
DARKRMS SC	Median of the dark rms (detector noise) in the Science detector [adu]

MEANDARK FT	Mean of the dark level in the FT detector [adu]
DARKRMS FT	Median of the dark rms (detector noise) in the Fringe-Tracker detector [adu]

7.3 gravity_p2vm

This recipe reduces the internal calibrations. As a special sequence of shutter opening is required, it is advised to always build the SOF with a complete sequence of files obtained within a single execution of the p2vm calibration template. However it is still possible to input a SOF with DARK_RAW only, or DARK_RAW and FLAT_RAW only. It is also possible to input a SOF with some already processed calibration (e.g WAVE).

1. Compute the dark, write product
2. Compute the flat, write product
3. Compute the badpixels, write product
4. Compute the spectral calibration, write product
5. Compute the p2vm, write product

Input

DO.CATG	
DARK_RAW	raw dark, all shutters closed (DPR.TYPE=DARK)
FLAT_RAW x4	raw flats, one shutter open (DPR.TYPE=FLAT)
P2VM_RAW x6	raw p2vms, two shutters open (DPR.TYPE=P2VM)
WAVE_RAW	raw wavelength calibration for FT (DPR.TYPE=WAVE)
WAVESC_RAW	raw wavelength calibration for SC (DPR.TYPE=WAVESC)

Output

PRO.CATG	
DARK	dark calibration
FLAT	flat calibration
BAD	badpixel calibration
WAVE	wave calibration
P2VM	p2vm calibration

Parameters

Name	short description
–static-name	Use static names for the products (for ESO). [FALSE]

–debug-file	Save additional debug file(s). [FALSE]
–preproc-file	Save the PREPROC intermediate product. [FALSE]
–bias-method	Methode to average the biaspixels when cleaning-up the SC detector (only apply to MED and LOW). Ideally the same value shall be used when reducing the DARK with gravity_dark and the OBJECT with gravity_vis. <MEDIAN | MEDIAN_PER_COLUMN> [MEDIAN]
–bad-dark-threshold	the rms factor for dark bad pixel threshold. [10]
–profile-mode	Methode to compute the profile.PROFILE corresponds to the pixel intensities measured in the FLAT files (Gaussian like with FWHM of approx 1.5 pixel). This is the AUTO for the Low and Med spectral resolution. GAUSS corresponds to a Gaussian fit of the pixel intensities measured in the FLAT files. BOX corresponds to a box-card of 6 pixels centered on the spectra measured in the FLAT files. This is the AUTO for High spectral resolution. <AUTO | PROFILE | GAUSS | BOX> [AUTO]
–force-badpix-to-zero	Force the badpixel to zero in profile. [TRUE]
–profile-width	Width the detector window extracted around the default position of each spectra, and on which the profile will be applied to perform the extraction. [6]
–force-wave-ft-equal	Force the spatial order of the wavelength 2D fit for FT to zero. [FALSE]
–phase-calibration	This option changes the phase reference of the P2VM: NONE defines phiA(lbd) at zero for all baselines (P2VM calibrates only the internal phase-shift of the beam combiner); CLOSURE defines phiA(lbd) at zero for baselines 01, 02 and 03 (P2VM calibrates the phase-shift and the closure-phase of the beam combiner); DISP defines phiA(lbd) to have zero mean and minimum GD for baselines (01,02,03); (P2VM calibrates the phase-shift, the closure-phase and the spectral-dispersion of the beam combiner); FULL defines phiA(lbd) to have zero-GD for baselines (01,02,03). <NONE | CLOSURE | DISP | FULL> [CLOSURE]

Quality control

QC in BAD	short description
BADPIX ACQ	Total number of bad pixels on the ACQ detector
BADPIX SC	Total number of bad pixels on the SC detector
BADPIX FT	Total number of bad pixels on the FT detector
BADPIX_DARK SC/FT	Pixels with weird mean level
BADPIX_RMS SC/FT	Pixels with weird noise level
BADPIX_FLAT SC/FT	Pixels non-responding to illumination

QC in FLAT	short description
PROFILE_CENTER SC1 MED	[pixel] position of the first spectra on SC detector
PROFILE_WIDTH SC1 MED	[pixel] width of the first spectra on SC detector
PROFILE_CENTER SC13 MED	[pixel] position of the 13d spectra on SC detector
PROFILE_WIDTH SC13 MED	[pixel] width of the 13d spectra on SC detector
MEANGAIN SC	Mean gain [ADU/e] for SC detector
MEANGAIN FT	Mean gain [ADU/e] for FT detector
QC in P2VM	short description
FLUX_SCi AVG	[e/DIT/chanel/output] flux in SC (mean of files)
FLUX_FTi AVG	[e/DIT/chanel/output] flux in FT (mean of files)
P2VM_COHERENCE_AVG_SC	Average instrumental contrast of SC
P2VM_COHERENCE_AVG_FT	Average instrumental contrast of FT
P2VM_COHERENCE_SCij	Average instrumental contrast of SC for pair ij
P2VM_COHERENCE_FTij	Average instrumental contrast of FT for pair ij
QC in WAVE	short description
REFWAVE1	Reference wavelengh [m] for the below parameters
REFPOS1 SCi	Position [pix] of the REFWAVE1 in output SCi
REFPOS1 FTi	Position [pix] of the REFWAVE1 in output FTi
REFWAVE2	Reference wavelengh [m] for the below parameters
REFPOS2 SCi	Position [pix] of the REFWAVE2 in output SCi
REFPOS2 FTi	Position [pix] of the REFWAVE2 in output FTi
WAVE_CORR	Model to convert the glass wavelength in vacuum wavelength
WAVE_CORR N0	Paramater of above model
WAVE_CORR N1	Paramater of above model
WAVE_CORR N2	Paramater of above model
MINWAVE SC/FT	Min wavelength [m] of SC/FT channels
MAXWAVE SC/FT	Max wavelength [m] of SC/FT channels
RMSWAVE SC/FT	Rms of residuals during polynomial wavelength fit

7.4 gravity_eop

7.5 gravity_eop

This recipe downloads the latest version of the Earth Orientation Parameter and DUT from the IERS site. File is created in the current directory. A web connection is required.

1. Download the IERS data
2. Convert into CPL table
3. Write product

Input

DO.CATG	
None	No input

Output

PRO.CATG	
EOP_PARAM	EOP calibration file (gravity_eop_calib.fits)

Parameters

Name	short description
–eop_host	FTP Host to retrieve the EOP from. [ftp.iers.org]
–eop_urlpath	FTP URL path of the EOP file to retrieve. [/products/eop/rapid/standard/finals2000A.data]

7.6 gravity_vis

This recipe is associated to the observations template. Its reduces the raw data acquired on calibrator or science targets and computes the uncalibrated visibilities, saved in an OIFITS file. If several OBJECT are provided, the recipe will reduce all of them and merge the resulting data into a single OIFITS. If several SKY_RAW are provided, the recipe reduces the first OBJECT with the first SKY file. Then each new OBJECT with the next SKY. When the number of SKYs is reached, the recipe loops back to first SKY file (so if the number of SKYs is larger than the number of OBJECTs, the last SKY won't be used). The recipe will reduce the data even if no SKY or no DARK is provided. However this will lead to wrong estimate of the visibility and squared visibility of the object. If the file DIAMETER_CAT is not provided, the recipe will use the diameter provided in the header to compute the transfer function QC parameters. The tag in the DO category can be SINGLE/DUAL and CAL/SCI. They should reflect the instrument mode (SINGLE or DUAL) and the DPR.CATG of the observation (SCIENCE or CALIB). The tag in the PRO.CATG category will be SINGLE/DUAL and CAL/SCI depending on the input tag.

1. Load the input file (loop on input OBJECT files)
2. Extract the spectrums (use BAD, DARK, SKY, FLAT files)
3. Interpolate the spectrums into a common wavelength table (use WAVE file)
4. Compute the real-time visibilities (use P2VM file)
5. Compute additional real-time signals (SNR, GDELAY...)
6. Compute selection flags (= flag frames with SNR lower than threshold, vFactor lower than threshold...)
7. Average the real-time visibilities, considering the selection flag
8. Write the product

Input

DO.CATG	
FLAT	flat calibration (PRO.CATG=FLAT)
BAD	badpixel calibration (PRO.CATG=BAD)
WAVE	wave calibration (PRO.CATG=WAVE)
P2VM	p2vm calibration (PRO.CATG=P2VM)
DARK	dark calibration (PRO.CATG=DARK)
SINGLE_SCI_RAW	raw object (DPR.TYPE=OBJECT,SINGLE)
SINGLE_SKY_RAW	raw sky (DPR.TYPE=SKY,SINGLE)
DISP_MODEL (opt)	fiber dispersion model (PRO.CATG=DISP_MODEL)
DIAMETER_CAT (opt)	catalog of diameter (PRO.CATG=DIAMETER_CAT)

Output

PRO.CATG	
SINGLE_SCI_VIS	OIFITS file with uncalibrated visibilities
SINGLE_SKY (opt)	sky map
SINGLE_SCI_P2VMRED (opt)	intermediate product (see detailled description of data)
SPECTRUM (opt)	intermediate product (see detailled description of data)
PREPROC (opt)	intermediate product (see detailled description of data)

Parameters

Name	short description
–static-name	Use static names for the products (for ESO). [FALSE]
–bias-subtracted-file	Save the BIAS_SUBTRACTED intermediate product. [FALSE]
–spectrum-file	Save the SPECTRUM intermediate product. [FALSE]
–preproc-file	Save the PREPROC intermediate product. [FALSE]
–p2vmreduced-file	Save the P2VMRED intermediate product. [FALSE]
–astro-file	Save the ASTROREDUCED intermediate product. [FALSE]
–average-vis	Average the observation from the different input files (if any) in the output product, instead of simply appending them. [FALSE]
–bias-method	Methode to average the biaspixels when cleaning-up the SC detector (only apply to MED and LOW). Ideally the same value shall be used when reducing the DARK with gravity_dark and the OBJECT with gravity_vis. <MEDIAN | MEDIAN_PER_COLUMN> [MEDIAN]

<table>
<tr><td>–ditshift-sc</td><td>Shift the time of SC DITs by an integer value to account for lost frames in exposure (issue on the instrument side, report to instrument team). The time of all DITs in exposure are increased by ditshift x PERIOD, hence ditshift=+1 means the system has lost one DIT. ditshift can be 0, positive, or negative. [0]</td></tr>
<tr><td>–nsmooth-snr-ft</td><td>Number of sample to average coherently when computing the real-time SNR and GDELAY of the FT (shall corresponds to the atmospheric coherence time). The runing integration window is actually -nsmooth -> +nsmooth. [5]</td></tr>
<tr><td>–reference-phase-sc</td><td>Compute the reference phase of the SC from the FT (FT, normal mode) or a self-reference build from a fit of the SC phase itself for each DIT (SC). <SC | FT> [FT]</td></tr>
<tr><td>–snr-min-ft</td><td>SNR threshold to accept FT frames (>0). It rizes the first bit («0) of column REJECTION_FLAG of FT. [3.0]</td></tr>
<tr><td>–global-state-min-ft</td><td>Minimum OPDC state to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [2.0]</td></tr>
<tr><td>–global-state-max-ft</td><td>Maximum OPDC state to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [4.0]</td></tr>
<tr><td>–state-min-ft</td><td>Minimum OPDC state per baseline to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [1.0]</td></tr>
<tr><td>–tracking-min-sc</td><td>Minimum ratio of accepted FT frames to accept SC frames (0..1), that is, for each SC DIT, the fraction of the time the REJECTION_FLAG of the FT is not 0. It rizes the first bit («0) of column REJECTION_FLAG of SC. [0.8]</td></tr>
<tr><td>–vfactor-min-sc</td><td>vFactor threshold to accept SC frame (0..1). [0.1]</td></tr>
<tr><td>–debias-sc</td><td>Subtract the V2 bias from SC. [TRUE]</td></tr>
<tr><td>–debias-ft</td><td>Subtract the V2 bias from FT. [TRUE]</td></tr>
<tr><td>–nboot</td><td>Number of bootstrap to compute error (1..100). [20]</td></tr>
<tr><td>–vis-correction-sc</td><td>Correction of SC visibility from the losses du to long-exposure. Either using the measured visibility losses with the FT (VFACTOR and/or PFACTOR, see the manual for the algorithms) or by forcing the SC visibilities to match those of the FT (FORCE). Possible choices are. <VFACTOR | PFACTOR | VFACTOR_PFACTOR | FORCE | NONE> [VFACTOR]</td></tr>
<tr><td>–flat-flux</td><td>Normalize the flux (stored in OI_FLUX binary extension) with instrument transmission recorded in the nput P2VM calibration map. Consequently, the flux quantity is either the intensity level recorded n the detector, thus including the instrument transmission (FALSE); or the intensity level at the instrument entrance (TRUE). [FALSE]</td></tr>
</table>

–average-sky	Average the SKYs into a master SKY. If FALSE, the recipe loops over the SKY to reduce each OBJECT with a different SKY. [FALSE]
–reduce-acq-cam	If TRUE, reduced ACQ_CAM images. [FALSE]

Quality control

QC in VIS	short description
CHECK FLAGS	Number of messages about data integrity
CHECK MSGi	Message about data integrity
TRANS PROFILE SC	[e/e] numerical transmission at profile extraction
TRANS INTERP SC	[e/e] numerical transmission at interpolation
TRANS P2VM SC	[e/e] numerical transmission at P2VM extraction
PHASE_FTij RMS	[rad] phase residuals when tracking
TRACKING_RATIO_FTij	[%] ratio of tracking time
ACCEPTED_RATIO_FTij	[%] fraction of accepted FT frame in the final averaging
ACEEPTED_RATIO_SCij	[%] fraction of accepted SC frame in the final averaging
TAU0 OPDCij	[s] tau0 for variance of 1 rad^2, computed from the fringe-tracking command to the piezo actuator.
VFACTORij_P1 AVG	mean v-factor for polar 1
VISAMP_SCij_P1 AVG	[deg] mean VISAMP for Science Combiner for polar 1
...	see all parameters
GD_SCij_P1 AVG	[m] Mean Group-Delay on Science Combiner
FLUX_FTi_P1 AVG	[e/total_int_time] mean flux over channels for Fringe Tracker
FLUXRATE_FTi_P1 SUM	[e/s] sum over channels for Fringe Tracker

If the observation is of type CALIB, and if the DIAMETER_CAT was provided, the recipe also computes the following QC:

QC in VIS	short description
TF TRANS_FTij	Total transmission of FT
TF TRANS_SCij	Total transmission of SC
TF VISMOD_SCij RELERR	TF relative error from diameter error
TF VISAMP_SCij_P1 MED	TF median over channels
TF VIS2_SCij_P1 MED	TF median over channels

7.7 gravity_vis_from_p2vmred

This recipe averages the real-time data of P2VMRED files into a VIS product. It allows to run the reduction with different parameters (for instance for SNR thresholding) without having to re-reduce the files from scratch. Typically the reduction is 4x faster when started from this intermediate product. The tag in the DO category can be SINGLE/DUAL and CAL/SCI. They should reflect the mode (SINGLE or DUAL) and the DPR.CATG of the observation (SCIENCE or CALIB). The tag in the PRO.CATG category will be SINGLE/DUAL and CAL/SCI depending on the input tag.

1. Load the input file (loop on input files)
2. Update the selection flag
3. Average the real-time visibilities
4. Write the product

Input

DO.CATG	
SINGLE_SCI_P2VMRED	Input intermediate product

Output

PRO.CATG	
SINGLE_SCI_VIS	OIFITS with uncalibrated visibilities

Parameters

Name	short description
–static-name	Use static names for the products (for ESO). [FALSE]
–average-vis	Average the observation from the different input files (if any) in the output product, instead of simply appending them. [FALSE]
–nsmooth-snr-ft	Number of sample to average coherently when computing the real-time SNR and GDELAY of the FT (shall corresponds to the atmospheric coherence time). The runing integration window is actually -nsmooth -> +nsmooth. [5]
–snr-min-ft	SNR threshold to accept FT frames (>0). It rizes the first bit («0) of column REJECTION_FLAG of FT. [3.0]
–global-state-min-ft	Minimum OPDC state to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [2.0]
–global-state-max-ft	Maximum OPDC state to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [4.0]
–state-min-ft	Minimum OPDC state per baseline to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [1.0]
–tracking-min-sc	Minimum ratio of accepted FT frames to accept SC frames (0..1), that is, for each SC DIT, the fraction of the time the REJECTION_FLAG of the FT is not 0. It rizes the first bit («0) of column REJECTION_FLAG of SC. [0.8]

–vfactor-min-sc	vFactor threshold to accept SC frame (0..1). [0.1]
–debias-sc	Subtract the V2 bias from SC. [TRUE]
–debias-ft	Subtract the V2 bias from FT. [TRUE]
–nboot	Number of bootstrap to compute error (1..100). [20]
–vis-correction-sc	Correction of SC visibility from the losses du to long-exposure. Either using the measured visibility losses with the FT (VFACTOR and/or PFACTOR, see the manual for the algorithms) or by forcing the SC visibilities to match those of the FT (FORCE). Possible choices are. <VFACTOR | PFACTOR | VFACTOR_PFACTOR | FORCE | NONE> [VFACTOR]
–use-existing-rejection	Use existing rejection flags (ignore related options). [FALSE]

7.8 gravity_viscal

This recipe calibrates the visibilities acquired on science target using visibilities acquired on a calibrator target. If the DIAMETER_CAT is not provided, the recipe will use the diameter provided in the header to compute the transfer function QC parameters. The corresponding keywords are INS.SOBJ.DIAMETER and FT.ROBJ.DIAMETER. The OI_FLUX data are not yet calibrated. The tag in the DO category can be SINGLE/DUAL and CAL/SCI. They should reflect the mode (SINGLE or DUAL) and the DPR.CATG of the observation (SCIENCE or CALIB). The tag in the PRO.CATG category will be SINGLE/DUAL and CAL/SCI depending on the input tag.

1. Loop on all input CALIB files, compute the TF for each of them and write the corresponding product
2. Loop on all input SCIENCE files, interpolate the TF at that time, calibrate, and write the corresponding product

Input

DO.CATG	
SINGLE_SCI_VIS (>=1)	visibilities on sciences
SINGLE_CAL_VIS (>=1)	visibilities on calibrators
DIAMETER_CAT (opt)	catalog of stellar diameters

Output

PRO.CATG	
SINGLE_SCI_VIS_CALIBRATED	calibrated science visibilities
SINGLE_CAL_TF	Transfer Function (TF) estimated on calibrators
SINGLE_SCI_TF	TF interpolated at the time of sciences

Parameters

Name	short description
–static-name	Use static names for the products (for ESO). [FALSE]
–delta-time-calib . [3.6e+03]	Delta time to interpolate the TF [s]
–force-calib	Force the calibration, don't check setup. [FALSE]
–nsmooth-tfvis-sc	Smooth the TF spectrally by this number of spectral bin, to enhance SNR (only apply to VIS2, VISPHI, VISAMP, T3PHI, T3AMP). This parameter is ignored in spectral mode LOW. [0]
–nsmooth-tfflux-sc	Smooth the TF spectrally by this number of spectral bin, to enhance SNR (only apply to FLUX, RVIS, IVIS). This parameter is ignored in spectral mode LOW. [0]
–maxdeg-tfvis-sc	Fit the TF spectrally by a polynomial to enhance SNR (only apply to VIS2, VISPHI, VISAMP, T3PHI, T3AMP). This parameter is ignored in spectral mode LOW. [5]

7.9 gravity_postprocess

This recipe allows to manipulate the product of the GRAVITY pipeline, mostly the VIS. It permits to merge several files together into a single VIS file with all observations; to average the observations of one or several VIS file to increse the SNR; to remove some data (FT, SC); and to resample the SC observation with spectral binning. The list of input files can be P2VMRED, VIS, VIS_CALIBRATED (or even RAW for some parameters). However they should all be compatible in term of setup and observed objets !! Note that the recipe performs only litle checks of the input file content and structure. Thus the user shall ensure the input files are conformable (same polarisation and spectral mode for instante)

1. Load the files
2. Execute request from user
3. Write product

Input

DO.CATG	
Input files	see above

Output

PRO.CATG	
POSTPROCESSED	Output file

Parameters

Name	short description
–average-vis	Average the observation from the different input files (if any) in the output product, instead of simply appending them. [FALSE]
–fluxerr-sc	Force the uncertainty in FLUX of SC. [0.0]
–visamperr-sc	Force the uncertainty in VISAMP of SC. [0.0]
–visphierr-sc	Force the uncertainty in VISPHI of SC. [0.0]
–vis2err-sc	Force the uncertainty in VIS2 of SC. [0.0]
–force-merge	Force merging even if inconsistent data. [FALSE]
–remove-ft	Remove FT extensions. [FALSE]
–remove-sc	Remove SC extensions. [FALSE]
–remove-opdc	Remove OPDC extensions. [FALSE]
–remove-met	Remove METROLOGY related extensions. [FALSE]
–nbin-lambda-sc	Bin SC extensions in spectral dimension. [0]

7.10 gravity_wavelamp

This recipe is associated to the template gravity_wavelamp. It reduces the raw file obtained with the Argon lamp (WAVELAMP) and process it so that it can be used to calibrate the fiber dispersion (recipe gravity_disp).

1. Extract the spectrums of the Argon exposure
2. Interpolate the spectrums into a common wavelength table
3. Measure the wavelength position of known Argon lines
4. Write the product

Input

DO.CATG	
FLAT	flat calibration (PRO.CATG=FLAT)
BAD	badpixel calibration (PRO.CATG=BAD)
WAVE	wave calibration (PRO.CATG=WAVE)
P2VM	p2vm calibration (PRO.CATG=P2VM)
WAVELAMP_RAW	long exposure of Argon lamp
DARK_RAW	dark of Argon exposure

Output

PRO.CATG	

WAVELAMP	spectrum of Argon, with position of lines

Parameters

Name	short description
–static-name	Use static names for the products (for ESO). [FALSE]

7.11 gravity_disp

This recipe is associated to the template GRAVI_all_disp. It measures the phases obtained on the internal source at the position of the Argon lines and various positions (= fiber stretch) of the Fibered Differential Delay Lines (FDDL). It deduces the linearity model and the dispersion model of the differential delay lines. These models are stored as polynomials versus wavelength.

1. Reduce all the input DISP files (see gravity_vis), write each product
2. Compute the dispersion parameters from this entire dataset
3. Write product

Input

DO.CATG	
FLAT	flat calibration (PRO.CATG=FLAT)
BAD	badpixel calibration (PRO.CATG=BAD)
WAVE	wave calibration (PRO.CATG=WAVE)
P2VM	p2vm calibration (PRO.CATG=P2VM)
DARK	dark calibration (PRO.CATG=DARK)
WAVELAMP	spectrum of Argon, with position of lines
DISP_RAW (>50)	raw dispersion

Output

PRO.CATG	
DISP_VIS	intermediate product
DISP_MODEL	dispersion model of FDDL

Parameters

Name	short description

–static-name	Use static names for the products (for ESO). [FALSE]
–preproc-file	Save the PREPROC intermediate product. [FALSE]
–p2vmreduced-file	Save the P2VMRED intermediate product. [FALSE]
–vis-file	Save the VIS intermediate product. [FALSE]
–nsmooth-snr-ft	Number of sample to average coherently when computing the real-time SNR and GDELAY of the FT (shall corresponds to the atmospheric coherence time). The runing integration window is actually -nsmooth -> +nsmooth. [5]
–reference-phase-sc	Compute the reference phase of the SC from the FT (FT, normal mode) or a self-reference build from a fit of the SC phase itself for each DIT (SC). <SC | FT> [FT]
–snr-min-ft	SNR threshold to accept FT frames (>0). It rizes the first bit («0) of column REJECTION_FLAG of FT. [30.0]
–global-state-min-ft	Minimum OPDC state to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [2.0]
–global-state-max-ft	Maximum OPDC state to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [4.0]
–state-min-ft	Minimum OPDC state per baseline to accept FT frames (>=0) It rizes the second bit («1) of column REJECTION_FLAG of FT. [1.0]
–tracking-min-sc	Minimum ratio of accepted FT frames to accept SC frames (0..1), that is, for each SC DIT, the fraction of the time the REJECTION_FLAG of the FT is not 0. It rizes the first bit («0) of column REJECTION_FLAG of SC. [0.8]
–vfactor-min-sc	vFactor threshold to accept SC frame (0..1). [0.8]
–debias-sc	Subtract the V2 bias from SC. [TRUE]
–debias-ft	Subtract the V2 bias from FT. [TRUE]
–nboot	Number of bootstrap to compute error (1..100). [1]
–vis-correction-sc	Correction of SC visibility from the losses du to long-exposure. Either using the measured visibility losses with the FT (VFACTOR and/or PFACTOR, see the manual for the algorithms) or by forcing the SC visibilities to match those of the FT (FORCE). Possible choices are. <VFACTOR | PFACTOR | VFACTOR_PFACTOR | FORCE | NONE> [NONE]

8 Re-creating the Static Calibration

8.1 Dispersion model

The **DISP_MODEL** static calibration can be recomputed with the recipes **gravity_wavelamp** and **gravity_disp**. The principle is to accurately measure the interferometric phases obtained for various position of the FDDL (= various strechting of the fibers) at the wavelengths of known Argon lines. The following dedicated RAW data are required:

WAVELAMP is a spectrum of the internal argon lamp.

DISP are interferometric observations of the internal source for various position of the FDDL (= different strech of the fibers).

8.2 Earth Orientation Parameters

The **EOP_PARAM** static calibration can be recomputed with the recipe **gravity_retrieve_eop**, which shall query the IERS webpage to obtain the best estimate of the past and futur EOP.

9 Detailed description of the data content

9.1 Recommended tools to browse data

RAW files can be conveniently opened with the last version of *QFitsView* [2] from Thomas Ott.

OIFITS product files can be conveniently open with the last version of *OIFits Explorer* [1] from the JMMC.

The consortium has built a python script to visualise most of the GRAVITY data, *run_gravi_visual.py*.

9.2 Table structure common to all data

The INSNAME header keyword specifies the combiner to which a table refers to, and thus allows to cross-reference with other tables. The EXTVER keyword specifies the instance of a table repeated in the OIFITS file to allow fast search in the tables (see python FITS class for instance). They can take the values: GRAVITY_SC' (10), GRAVITY_SC_P1' (11), GRAVITY_SC_P2' (12), GRAVITY_FT' (20), GRAVITY_FT_P1' (21), GRAVITY_FT_P2' (22).

The polarisation 'P1' in output products correspond to the 'S' regions in IMAGING_DETECTOR tables. The polarisation 'P2' in output products correspond to the 'P' regions in IMAGING_DETECTOR tables.

The IMAGING_DETECTOR_SC and IMAGING_DETECTOR_FT tables store the detector configurations based on the VLTI interface control document [5].

The IMAGING_DATA_SC and IMAGING_DATA_FT extensions store the detector data. The SC data are stored as image list, while the FT data are stored as tables.

The OI_WAVELENGTH tables store the wavelength table following the OIFITS standard. They shall be associated to the SC or FT using the INSNAME or EXTVER keywords.

OIFITS tables storing quantities per-beam (OI_FLUX) have a total of NDIT x 4 rows (or NEXP x 4 rows for final product). In these tables, the four beams are always ordered following 12341234...

OIFITS tables storing quantities per-baseline (OI_VIS, OI_VIS2) have a total of NDIT x 6 rows (or NEXP x 6 rows for final product). In these tables, the six baselines are always ordered following 123456123456... The baselines 1 to 6 are always the pair of the beams 1-2,1-3,1-4,2-3,2-4,3-4.

9.3 P2VM product

The P2VM_SC and P2VM_FT tables contain the following columns:

REGNAME : Detector region name, to match the IMAGING_DATA table.

TRANSMISSION : For each region (= output of the combiner), a ntel $\times$ nwave image with the transmission of each input beam in this region. Since the combination scheme is pairwise, normaly only 2 rows of this image shall be non-zero.

COHERENCE : For each region (= output of the combiner), a nbase × nwave image with the instrumental visibility of each pair of input beam in this region. Since the combination scheme is pairwise, normally only one single rows of this image shall be non-zero.

PHASE : For each region (= output of the combiner), a nbase × nwave image with the instrumental phase in radian of each pair of input beam in this region. Since the combination scheme is pairwise, normaly only one single rows of this image shall be non-zero.

C_MATRIX :

The OI_FLUX tables store the internal instrument transmission, measured on the internal light, following the OIFITS standard, in arbitrary units. These tables are used to later calibrate the measured flux of the object using the spectral shape of the internal transmission.

9.4 *_VIS and *_TF products

The files with PRO.CATG=*_VIS and *_TF follow the OIFITS standard, version 1. All information can be found in [4].

The OIFITS tables contain one row per baseline and per corresponding RAW exposure, that is all the frames from the individual exposure are averaged together.

The OIFITS tables are associated to the SC or FT using the INSNAME (and/or EXTVER) keywords.

9.5 SPECTRUM, PREPROC products

These are intermediate products used to debug the pipeline.

Both files contain the extracted spectra of each output of the combiners. In SPECTRUM, the spectra are in pixel space, thus on different wavelength grids. In PREPROC, the spectra have been re-interpolated into a common wavelength grid.

Columns in the SPECTRUM_DATA tables

TIME [μs] : time of the frame, in [us], from the PRC.ACQ.START time from header (RMN recording start).

DATAi [e] : The spectrum of the flux from output i of the combiner.

DATAERRi [e] : The spectrum of the theoretical error of the flux from output i of the combiner, including detector and photonic variances.

9.6 *_P2VMRED products

The files with PRO.CATG=*_P2VMRED use elements of the OIFITS format [4], but are non-standard for the TIME colums. They also include many additional columns to store intermediate signal.

Beware that the OIFITS tables contain one row per frame: the number of row can be huge for the FT tables.

Columns in the OI_VIS table of the SC

TARGET_ID : id listed in OI_TARGET

TIME [μs] : time of the frame, in [us], from the PRC.ACQ.START time from header (RMN recording start).

MJD [day] :

INT_TIME [s] : integration time of this frame

VISDATA [e,e] : complex coherent flux of SC in this frame

VISERR [e,e] : theoretical complex error on the coherent flux

UCOORD [m] : uv-plane of this SC frame

VCOORD [m] : uv-plane of this SC frame

STA_INDEX : station index in the OI_ARRAY

PHASE_MET_FC [rad] : unwrapped FT-SC phase as computed by the DRS algorithm

PHASE_MET_TEL [rad] : unwrapped FT-SC phase as computed by the DRS algorithm, mean of 4 diodes

OPD_MET_FC [m] : unwrapped SC-FT delay as computed by the TAC algorihm

OPD_MET_TEL [m] : unwrapped SC-FT delay as computed by the TAC algorihm, 4 diodes

OPD_DISP [m] : spectra of OPD introduced by fibers, including dispersion.

VISDATA_FT [e,e] : <VISDATA> spectra of FT (integrated in this SC frame)

VISPOWER_FT [e**2] : <|VISDATA|**2> spectra of FT (integrated in this SC frame)

VISVAR_FT [e**2] : <|VISERR|**2> spectra of FT (integrated in this SC frame)

GDELAY_FT [m] : real-time GD computed from VISDATA_FT

PHASE_REF [rad] : reference phase, actually -1*argVISDATA_FT, re-interpolated in the SC wavelength.

PHASE_REF_COEFF [rad] : polynomial coefficients fit to argVISDATA_FT and used to extrapolate to the SC wavelengths, in units of $(\lambda - \lambda_{mean})/(\lambda_{max} - \lambda_{min})$

FIRST_FT : first FT frame in this SC frame

LAST_FT : last FT frame in this SC frame

FIRST_MET : first MET frame in this SC frame

LAST_MET : last MET frame in this SC frame

GDELAY [m] : real-time GD computed from VISDATA

SNR : real-time SNR

GDELAY_BOOT [m] : best GD estimate, taking into account closing triangles

SNR_BOOT : best SNR estimate, taking into account closing triangles

V_FACTOR_FT : measured visibility loss on the FT

V_FACTOR : predicted visibility loss of this SC frame (re-interpolation of V_FACTOR_FT on the SC wavelengths)

P_FACTOR : predicted visibility loss of this SC frame due to photometric flickering (based on the real-time photometry of the FT)

F1F2 Estimate of geometric flux of this frame.

FRINGEDET_RATIO : fraction of FT frame accepted in this SC frame

REJECTION_FLAG : this frame is accepted/rejected

E_U, E_V, E_W : Local celestial {u,v,w} (East, North, Toward observer = OIFITS standard) expressed in local terrestrial (East, North, Up) at Paranal reference. It is usefull for recomputing the projected baseline from physical baseline.

E_Az : Vector product of the E_W and the Zenith directions, expressed in the local terrestrial (East, North, Up) at Paranal reference. Sitting on telescope, looking at the target, E_Az points toward left in the horizontal plane.

E_Zd : Vector product of the E_W and E_Az, expressed in the local terrestrial (East, North, Up) at Paranal reference. Sitting on telescope, looking at the target, E_Az points toward Nadir in the plane perpendicular to pointing direction.

PUPIL_U [m] : lateral shift of pupil (in uv reference).

PUPIL_V [m] : lateral shift of pupil (in uv reference).

PUPIL_W [m] : focus shift of pupil (in uv reference).

Columns in the OI_VIS table of the FT

TARGET_PHASE [rad] : target phase of the loop, including the Sylvester modulation

STATE : baseline tracking state as reported by OPDC

GDELAY [m] : real-time GD computed from VISDATA

GDELAY_BOOT [m] : best GD estimate, accounting closing triangles

SNR : real-time SNR

SNR_BOOT : best SNR estimate, accounting closing triangles

REJECTION_FLAG : this frame is accepted/rejected

F1F2 Estimate of geometric flux of this frame.

PHASE_REF [rad] : self-reference phase.

Columns in the OI_FLUX table of the SC

TARGET_ID : id listed in OI_TARGET

TIME [us] : time of the frame, in [us], from the PRC.ACQ.START time from header (RMN recording start).

MJD [day] :

INT_TIME [s] : integration time of this frame

FLUX [e] : flux

FLUXERR [e] : theoretical error on flux

STA_INDEX : station index in the OI_ARRAY

TOTALFLUX_SC [e] : total flux of SC in this SC frame (integrated over spectrum)

TOTALFLUX_FT [e] : total flux of FT in this SC frame (integrated over spectrum)

OPD_MET_FC [m] : unwrap SC-FT delay as computed by the TAC algorihm

OPD_MET_TEL [m] : unwrap SC-FT delay as computed by the TAC algorihm, 4 diodes.

FT_POS [V] : mean FT FDDL strain gauge voltage during this frame

SC_POS [V] : mean SC FDDL strain gauge voltage during this frame

FDDL [m] : The mean of SC_FDDL and FT_FDDL, where these value are the FDDL strain gauge measurments, corrected from non-linearity and converted in [m].

FIRST_FT : first FT frame in this SC frame

LAST_FT : last FT frame in this SC frame

FIRST_MET : first MET frame in this SC frame

LAST_MET : last MET frame in this SC frame

FIRST_FDDL : first FDDL frame in this SC frame

LAST_FDDL : last FDDL frame in this SC frame

Columns in the OI_FLUX table of the FT

STATE : telescope state as reported by OPDC

Columns in the OI_VIS_MET table

The table has the same structure as the OI_FLUX table, that is one row per beam (12341234124...), and thus nsample x 4 rows.

PHASE_FC [rad] : phases at combiner, unwrap by pipeline algorithm (FT-SC)

PHASE_TEL [rad] : 4 diodes phases at telescope, unwrap by pipeline algorithm (FT-SC)

OPD_FC [m] : OPD at telescope, unwrap by TAC algorithm (SC-FT)

OPD_TEL [m] : 4 diodes OPD at telescope, unwrap by TAC algorithm (SC-FT)

FLAG_FC, FLAG_TEL : flags computed by TAC algorithm

VAMP_FC_FT , VAMP_FC_SC, VAMP_TEL_FT, VAMP_TEL_SC: Volt amplitude

Columns in the OI_VIS_ACQ table

The table has the same structure as the OI_FLUX table, that is one row per beam (12341234124...), and thus nsample x 4 rows.

TIME [us] : time of the frame, in [us], from the PRC.ACQ.START

PUPIL_NSPOT : number of spot detected in the pupil sensor (maximum is 16 = 4 diodes x 4 sub-apertures).

PUPIL_X [pix] : horizontal shift of pupil (in detector).

PUPIL_Y [pix] : vertical shift of pupil (in detector).

PUPIL_Z [pix] : focus shift of pupil (in detector).

PUPIL_R [deg] : rotation of pupil diode (in detector).

PUPIL_U [m] : lateral shift of pupil (in uv reference).

PUPIL_V [m] : lateral shift of pupil (in uv reference).

PUPIL_W [m] : focus shift of pupil (in uv reference).

Columns in the OPDC table

TIME [us] : time of the frame, in [us], from the PRC.ACQ.START time from header (RMN recording start).

PIEZO_DL_OFFSET

VLTI_DL_OFFSET

STATE : global fringe tracking state

BASELINE_STATE : fringe tracking state per telescope and baseline (scrambled)

STEPS : target phase modulation per baseline (scrambled), in units of pi/8

9.7 ASTRORED product

This is a light version of the P2VMRED file. It is meant to develop the astrometric mode of GRAVITY.

9.8 DISP_MODEL

The data are stored in the table DISP_MODEL. There are one line per GRAVITY beam. The columns are:

BETA : Mean optical index of the SC and FT fibers. The vector contains the β_i coefficients such that the index at wavelength λ is given by $n(\lambda)/n(\lambda_{MET}) = \sum_i a_i\,(1/\lambda - 1/2.2)^i$ (with λ in [um]).

GAMMA : Differential optical index between the SC and the FT fibers. The vector contains the γ_i coefficients such that the differential index at wavelength λ is given by contains the β_i coefficients such that the index at wavelength λ is given by ??

LIN_FDDL_FT : Linearity coefficients to convert the FDDL signal in [V] to stretching length in [m]. The vector contains the a_i coefficients such that $L(V) = \sum_i a_i\,V^i$.

LIN_FDDL_SC : Same for the SC.

10 Algorithms

In the following, f is the index of individual DIT (that is from 0 to NDIT-1, whose typicall values are 30 for SC and 300000 for FT), ij are the pixel indices on the detector (i being the spatial direction, and j being the spectral direction), l is for the spectral channel, o is the combiner output (so called region in the code, [0..23] or [0..47] if polarisations are split), b is the baseline [0..5], and t is the telescope/beam [0..4].

10.1 Correction of detector bias

The SC frames are corrected for the detector bias using the value of some specific pixels, supposedly non-illuminated. In LOW and MED spectral mode, the bias per frame is estimated as the mean over all the pixels of the *bias lines* interleaved between the spectra of each region.

In HIGH spectral mode, the bias per frame is estimated as the mean over all pixles of the *bias columns* at the edge of the detector.

In both case, there is thus a single, scalar bias value for the entire frame for each frame.

10.2 Spectrum extraction

The implemented spectrum extraction Y_{foj} from the 2D image X_{fij} is based on a profile image p_{oij}.

Profile definition

In LOW and MED spectral modes, the profile is identical to the one observed with the sequences of FLAT files. It resembles a Gaussian function with FWHM of 1.5 pixel. To ensure the overall flux is conserved in the extraction, we apply the following normalization which assumes the shape of the object spectrum is perfectly matched by the profile itself:

$$p_{oij} = p_{oij} \cdot \frac{\sum_i p_{oij}}{\sum_i p_{oij}^2} \tag{1}$$

In HIGH mode, the profile is a boxcar of 6 pixels around the center of the best-fit Gaussian on the observed profile in the FLAT files. To ensure flux conservation, this boxcar is either 0 (outside) or 1 (inside).

Bad pixels in profile

The bad pixels are forced to zero in the profile. Consequently, a profile with bad pixels will lead to a reduced amount of detected flux. This effect is calibrated by the P2VM algorithm because the P2VM coefficients and the data are affected by the same amout of flux losses. A worst, for some spectral channels, the spectra of one output (e.g A) can be forced zero if all the pixels are bad. The P2VM then relies on the remaining BCD outputs only.

Extracted spectrum

$$Y_{foj} = g \sum_i (X_{fij} - S_{ij})\, p_{oij} \tag{2}$$

where $S_{i,j}$ is the mean image measured on the SKY. The sum $\sum$ is obviously performed in the compressed spatial direction.

Finally, g is the conversion gain from [ADU] to [e].

Variance of the extracted spectrum

Introducing the photon and detector noise, the variance of the extracted spectrum is:

$$\mathrm{var}(Y)_{foj} = g \sum_i (X_{fij} - D_{ij})\, p_{oij}^2 \;+\; g^2 \sum_i \sigma_{ij}^2\, p_{oij}^2 \tag{3}$$

where $D_{i,j}$ is the mean level measured on the DARK, and σ_{ij}^2 is the variance measured on the DARK.

10.3 Re-interpolation to a common wavelength

Modified target wavelength for SC

For the output (region) o, let's call j_{ol} the sample just below the target wavelength λ_l, and $j_{ol}+1$ the sample just above. Of course j depends on the region because their wavelength tables are different.

For the SC, this target wavelength is slightly modified, in a different way, for each region:

$$\lambda'_{ol} = \lambda_{o\,j_{ol}} + \frac{(\lambda_l - \lambda_{o\,j_{ol}})\,(\lambda_{o\,j_{ol}+1} - \lambda_{o\,j_{ol}})}{(\lambda_l - \lambda_{o\,j_{ol}}) + (\lambda_{o\,j_{ol}+1} - \lambda_l)\,.\,\frac{F_{o\,j_{ol}+1}}{F_{o,j_{ol}}}} \tag{4}$$

where F_{oj} is the flat measured on the internal light, extracted the same way as the data. This modification ensures that we later interpolate to a common *effective wavelength* for all regions. It also ensures that spectral channels whose interpolation includes a bad-pixel (forced to zero) are all set to zero.

Modified target wavelength for FT

For the FT, we don't modify the target wavelength:

$$\lambda'_{ol} = \lambda_l \tag{5}$$

Interpolation of flux and variance

The following coefficient a_{ol}

$$a_{ol} = \frac{\lambda_{o\,j+1} - \lambda'_{ol}}{\lambda_{o\,j+1} - \lambda_{o\,j}} \tag{6}$$

allows to linearly interpolate the fluxes:

$$Y_{fol} = a_{ol}\, Y_{fo\, j_{ol}} \quad + \quad (1 - a_{ol})\, Y_{fo\, j_{ol}+1} \tag{7}$$

and the variances:

$$\text{var}(Y)_{fol} = a_{ol}^2\, \text{var}(Y)_{fo\, j_{ol}} \quad + \quad (1 - a_{ol})^2\, \text{var}(Y)_{fo\, j_{ol}+1} \tag{8}$$

10.4 Extraction of the coherent fluxes and telescope fluxes via P2VM

The fluxes of each telescope F_{ftl} and the complex coherent flux of each base $R_{fbl} + iI_{fbl}$ are extracted from a matricial analysis of the profiles, based on the P2VM calibration

$$(F_{ftl}, R_{fbl}, I_{fbl}) = P2VM^o_{b/tl} \times Y_{fol} \tag{9}$$

The variances are propagated assuming no correlation between the input Y_{fol}.

$$(\text{var}(F)_{ftl}, \text{var}(R)_{fbl}, \text{var}(I)_{fbl}) = (P2VM^o_{b/tl})^2 \times \text{var}(Y)_{fol} \tag{10}$$

The $P2VM$ is a well conditioned matrix thanks to the design of the integrated beam combiner. To demonstra the underlying reasoning, let's consider a perfect $P2VM$. The four regions related to baseline b (say $0, 1, 2, 3$, also called ABCD regions of baseline b) can be combined together to build the following quantities:

- $R_{fbl} = Y_{f0l} - Y_{f2l}$,
- $I_{fbl} = Y_{f1l} - Y_{f3l}$ and
- $F_{ft_1l} + F_{ft_2l} = Y_{f0l} + Y_{f1l} + Y_{f2l} + Y_{f3l}$.

The two first are directly the complex coherent flux of baseline b, while the latter, combined with the constraints of the 5 other baselines, easily solve for the flux F_{ftl} of individual beams. In practice, the actual $P2VM$ matrix takes into account the exact interferometric phase-shift between the four ABCD regions, and the relative photometric throughputs.

10.5 Computation of SNR

Individual SNR

The Signal to Noise Ratio (SNR) of each baseline and each frame of the the FT is computed using a running mean of the complex coherent flux over 10 consecutive samples. The complex coherent flux of the two polarisations, if any, are also averaged together after having recentered them to a common mean phase.

$$SNR_{fb} = \frac{(\sum_{f_r\, l} R_{f_r\, bl})^2 + (\sum_{f_r\, l} I_{f_r\, bl})^2}{\sum_{f_r\, l} \text{var}(R)_{f_r\, bl} + \sum_{f_r\, l} \text{var}(I)_{f_r\, bl}} \tag{11}$$

where f_r is the f index running in the interval $\in \{f - 5\ ,\ f + 5\}$ in order to implement the smoothing over 10 samples.

Bootstrapped SNR

If fringes are detected on two consecutive baselines, then, by nature of light, fringes are detected on the closing baseline. The purpose of bootstrapping is to increase the confidence on the signal for the baselines with lowest SNR (whose SNR may well be null for astrophysical reason), by looking at these closing baseline.

A "bootstraped" SNR is computed for each baseline and each frame, as the maximum between the SNR of this baseline and all closing triangles. For instance for the baseline $b = b_{ij}$ among the beams i, j, k, l:

$$SNRB_{fb_{ij}} = max\{\ SNR_{fb_{ij}},\ min\{SNR_{fb_{ik}}\ ,\ SNR_{fb_{kj}}\}, min\{SNR_{fb_{il}}\ ,\ SNR_{fb_{lj}}\}\ \} \quad (12)$$

The quantities SNR and SNRB are also computed for the SC, although they are not used in the processing.

10.6 Computing the vFactor

The purpose of the vFactor is to estimate the visibility loss of each individual SC frame due the phase jittering, from an analysis of the FT real-time data. This visibility loss is derived as the ratio between the coherent integration (squared norm of complex sum over DITs) and the incoherent integration (sum over DITs of complex squared norm) of the complex coherent flux of FT, across each SC DIT.

A white-light vFactor is first computed for each SC frame f with the FT data:

$$v_{fb} = \frac{(\sum_{f_{FT}l} R_{f_{FT}bl})^2 + (\sum_{f_{FT}l} I_{f_{FT}bl})^2 - \sum_{f_{FT}l} \mathrm{var}(R)_{f_{FT}bl} - \sum_{f_{FT}l} \mathrm{var}(I)_{f_{FT}bl}}{\sum_{f_{FT}} (\sum_l R_{f_{FT}bl})^2 + \sum_{f_{FT}} (\sum_l I_{f_{FT}bl})^2 - \sum_{f_{FT}l} \mathrm{var}(R)_{f_{FT}bl} - \sum_{f_{FT}l} \mathrm{var}(I)_{f_{FT}bl}} \times \frac{1}{\mathrm{n}_{f_{FT}}} \quad (13)$$

where the sum over l is over the 6 spectral channels of the FT and the sum over f_{FT} is over the FT frames acquired *during* the corresponding SC frame. This white-light vFactor at λ_0 (the central wavelength of the FT) is then extrapolated to the SC channels with:

$$\widetilde{v}_{fbl} = \exp(-\ln(v_{fb})\,\frac{\lambda_0^2}{\lambda_l^2}\,) \quad (14)$$

This vFactor correction is proved to be very efficient as long as the FT astrophysical visibility remain larger than 0.1. For fully resolved baselines, the vFactor results into a indefinit 0/0 ratio. In such a sitation, and if on-axis, one could simply rescale the SC visibilities to the one of the FT (see options of the recipes). In off-axis case with a fully resolved object on the FT, there is not much to be done however.

10.7 Computing the pFactor

The purpose of the vFactor is to estimate the visibility loss of each individual SC frame due to flux flickering, from an analysis of the FT real-time flux data. This visibility loss is derived as the ratio between the sum over DITs of the geometry mean, and the geometric mean of the sum over DITs of the photometric flux of FT, across each SC DIT.

We compute a white-light pFactor with the following formula:

$$p_{bf} = \frac{\left[\sum_{f_{FT}} \sqrt{(\sum_l \mathrm{flux}_{f_{FT}t_1\,l})(\sum_l \mathrm{flux}_{f_{FT}t_2\,l})}\,\right]^2}{(\sum_{f_{FT}l} \mathrm{flux}_{f_{FT}t_1\,l})(\sum_{f_{FT}l} \mathrm{flux}_{f_{FT}t_2\,l})} \tag{15}$$

This pFactor is computed in the P2VMRED product but not used so far.

10.8 Frame rejection

A FT frame is rejected if any of the conditions are met:

- its bootstrapped $SNRB$ is below the threshold.
- the OPDC state of this baseline is below the threshold. The OPDC states are: 1 = IDEL, 2 = GD_TRACKING, 3 = PHASE_TRACKING, 4 = SEARCHING, 5 += internal calibrations.

A SC frame is rejected if any of the conditions are met:

- the fraction of accepted FT frame during this SC frame is below the threshold.
- its computed vFactor is below the treshold.

The frame selection is done on a per-baseline basis. That is the baselines have a different selection map, and thus will have a different effective time after the averaging process.

10.9 Phase referencing

Referencing the SC phases with the FT

The phase reference of a SC frame is the mean phase of the FT during this SC frame:

$$\Upsilon_{fbl} = \arctan(\sum_{f_{FT}} I_{f_{FT}bl}\ ,\ \sum_{f_{FT}} R_{f_{FT}bl}) \tag{16}$$

where the sum over f_{FT} is over the FT frames acquired during the corresponding SC frame.

This Υ, which have 6 spectral channels only, is interpolated/extrapolated into the wavelengths of the SC with a polynomial fit of order 2 (after properly unwrapping the phase along the spectral direction).

Self-referencing the FT phase

The phase reference of a FT frames is the runnning mean phase of the FT itself over few samples:

$$\Upsilon_{fbl} = \arctan(\sum_{f_r} I_{f_r\,bl}\ ,\ \sum_{f_r} R_{f_r\,bl}) \tag{17}$$

where f_r is in the interval $\in \{f-3\ ,\ r+3\}$, exluding $f_r = f$ to avoid biases. Note that this phase is not unwrapped neither temporally nor spectrally (SNR too low, and FT supposed to be near constant phase all the time thanks to real-time tracking).

10.10 Geometric flux

The geometric flux of a baseline is the product of the photometric fluxes of the two beams of this baseline. It is needed to then normalise the complex coherent fluxes into complex visibilities.

For the FT

The FT real-time photometry is affected by large photometric fluctuations, which make it going near zero (flux loss event) or even going to negative values because of noise. Since the geometric flux will enter the visibility normalisation in the denominator, we have to avoid as much as possible these near-zero events.

The photometric fluctuations are highly corrected in the spectral direction. On the other side, the important property of the fluctuations to be extracted are the temporary correlations. Therefore we first compute a broad-band mean of the real-time flux of each beam, that we also temporally smoothed:

$$F'_{ft} = \sum_{f_r\, l} F_{f_r\, tl} \tag{18}$$

where f_r is in the interval $\in \{f-5\ ,\ f+5\}$ (time smoothing). We then use this quasi real-time scaling to weight the mean spectrum. Only then the geometric mean is computed.

$$FF_{fbl} = (\sum_{f} F_{f\, t_1\, l})\, F'_{ft_1} \ \times\ (\sum_{f} F_{f\, t_2\, l})\, F'_{ft_2} \tag{19}$$

For the SC

For the science, the geometric flux is simply computed as:

$$FF_{fbl} = F_{f\, t_1\, l} \ \times\ F_{f\, t_2\, l} \tag{20}$$

10.11 Averaged flux estimator

The previous sections describe how the real-time quantities are extracted from every single FT and SC frame. We here describe the process of averaging these real-time quantities into final product. For the flux, all frames are simply co-added. There is no frame selection at all.

$$\widetilde{\mathrm{flux}}_{tl} = \sum_{f} F_{ftl} \tag{21}$$

Hence the final flux $\widetilde{\mathrm{flux}}_{tl}$ is the *sum* of all electron-events collected during the entire exposure and across all regions, for the beam t at channel l.

10.12 Averaged complex visibility estimator

For a given baseline, the averaged is computed only with accepted frames of this baseline.

The complex coherent flux vector is first rotated with the computed reference phase:

$$R'_{fbl} = \cos(\Upsilon_{fbl})\, R_{fbl} \;-\; \sin(\Upsilon_{fbl})\, I_{fbl} \tag{22}$$

$$I'_{fbl} = \sin(\Upsilon_{fbl})\, R_{fbl} \;+\; \sin(\Upsilon_{fbl})\, I_{fbl} \tag{23}$$

The visibilities of each frame are averaged together accounting for the visibility loss expected from the vFactor (only for SC, that is $v_{fbl} = 1.0$ for FT). Note that the coherent flux and the photometric flux are actually averaged first, before normalisation. The noise property is better than simply averaging the real-time normalised visibilities (Cauchy statistic).

Visibility amplitude

$$\widetilde{\mathrm{visAmp}}_{bl} = \frac{\sqrt{(\sum_f R'_{fbl})^2 + (\sum_f I'_{fbl})^2}}{\sum_f \sqrt{FF_{fbl}\; v_{fbl}}} \tag{24}$$

Visibility phase

$$\widetilde{\mathrm{visPhi}}_{bl} = \arctan(\sum_f I'_{fbl}\;,\; \sum_f R'_{fbl}) \tag{25}$$

The mean spectral slope (stored in the GDELAY_ASTROMETRY quantity) and mean spectral value (stored in the PHASE_ASTROMETRY quantity) are removed from the VISPHI quantity.

10.13 Average squared visibility estimator

The square visibilities of each frame are averaged together accounting for the visibility loss expected from the vFactor (only for SC, that is $v_{fbl} = 1.0$ for FT).

For a given baseline, the averaging is performed only with accepted frames of this baseline.

$$\widetilde{\mathrm{vis2}}_{bl} = \frac{\sum_f R^2_{fbl} + \sum_f I^2_{fbl} - \sum_f \mathrm{var}(R)_{fbl} - \sum_f \mathrm{var}(I)_{fbl}}{\sum_f (FF_{fbl}\; v_{fbl})} \tag{26}$$

10.14 Average closure-phase estimator

The averaged bispectrum of triplet b_{ijk} is computed as the coherent integration of the bispectrum of each frame:

$$\widetilde{B}_{b_{ijk}\,l} = \sum_f (R_{fb_{ij}l} + i\, I_{fb_{ij}l}).(R_{fb_{jk}l} + i\, I_{fb_{jk}l}).(R_{fb_{ik}l} - i\, I_{fb_{ik}l}) \tag{27}$$

For a given triplet, the integration is performed only with frames for which *all* three baselines forming the triangles are accepted.

Bispectrum phase

The closure phase is computed:

$$\widetilde{\mathrm{t3Phi}}_{b_{ijk}l} = \arctan(\widetilde{B}_{b_{ijk}\,l}) \tag{28}$$

Bispectrum amplitude

The bispectrum amplitude is also computed. However this quantity has not been verified yet.

$$\widetilde{\mathrm{t3Amp}}_{bl} = \frac{\widetilde{B}_{b_{ijk}\,l}}{\sum_f (F_{ft_i l}\; F_{ft_j l}\; F_{ft_k l}\; \sqrt{v_{fb_{ij}l}\, v_{fb_{ik}l}\, v_{fb_{jk}l}})} \tag{29}$$

10.15 Uncertainty on average quantities

The uncertainty on average quantities is computed by bootstrapping over the accepted frames. Basic fundamentals about this technic can be in the numerical recipes book, section 15.6.
See also: https://en.wikipedia.org/wiki/Bootstrapping_(statistics)

The implementation of the bootstraping method makes use of *segmentation* when the number of frame is larger than 100, or of *Monte-Carlo* when the number of frame is smaller than 5.

Segmentation (mostly for FT)

When the number of frames is larger than 100, the dataset is first split into a smaller number of *segments* (typically 20 to 100). First integration is done inside the segments. The final variance is estimated by bootstrapping over the segments. The number of segments does change the temporal sampling of the bootstrap, and thus the estimation of the uncertainties. We choose the number of segment so that each is about 1s length.

Note that this is mostly relevant for the FT, since the SC has often less than 100 frames (so each frame is a segment).

Monte-Carlo (mostly for very long DIT on SC)

When the number of frames is smaller than 5, the bootstrap method provides unrealistically small uncertainties. Hence the dataset is complemented with few fake frames (up to 5), on which a random realisation of the theoretical noise is added, using the theoretical variance.

This noise is added on the correlated flux quantity only, not on the photometric fluxes, vFactor, reference phase... Moreover it is clear that Monte-Carlo propagates the *fundamental* uncertainties only (photon and detector noise), but not the *atmospheric* noises (effect of injection fluctuation, tracking quality fluctuation).

As a matter of fact, the final uncertainty when the number of frames is very low can still be under-estimated.

10.16 Processing of MET and FDDL

From Volts to real time SC-FT phases

The signals of all diodes are analysed independently, without prior correction of the voltage, and converted into a real-time SC-FT phase for each diode. This real-time SC-FT phase is unwrapped and forced to match the OCS.MET reference. Both the unwrapping and the absolute reference are done with fringe-integer corrections only. These quantities are labeled OPD_TEL and OPD_FC in the OI_VIS_MET table.

Averaging SC-FT phase inside SC DIT

The phase of a diode is averaged inside each SC DIT, as a scalar quantity (not phasor). This averaged signal is stored for each beam in OI_FLUX table and each base in OI_VIS table. These quantities are labeled OPD_MET_TEL and OPD_MET_FC.

Average the FDDL inside SC DIT

The mean SC and FT FDDLs strain gauge voltage during each frame of the SC, and per beam, is stored in the OI_FLUX table. Columns are labelled SC_POS and FT_POS.

10.17 Applying dispersion correction to MET

Correction of FFDL non linearity

The DISP_MODEL provides, for each beam t:

- linSC_{tm} : the non-linearity coefficients of order m of the SC FDDL.
- linFT_{tm} : the non-linearity coefficients of order m of the FT FDDL.

SC_POS and FT_POS are first corrected from non-linearity of the strain gauge, and then co-added:

$$\mathrm{FDDL}_{tf} = \sum_m (\;\mathrm{linSC}_{tm}\,\mathrm{SC\_POS}^m_{tf} \;+\; \mathrm{linFT}_{tm}\,\mathrm{FT\_POS}^m_{tf} \;) \qquad (30)$$

Dispersion-included metrology signal

The DISP_MODEL allows to compute the *mean* refractive index of SC and FT FDDL at wavelength l for each beam t, normalized to the one at the metrology wavelength (hereafter called nmean_{tl}); and the *differential* refractive index between SC and FT FDDL wavelength l for each beam t, normalized to the one at the metrology wavelength (hereafter called ndiff_{tl}).

In practice, these two quantities are stored as a polynomial model versus wavenumber, centered in the middle of the K-band ($\lambda_0 = 2.2\mu$m):

$$\mathrm{nmean}_{tl} = \sum_m (\ \mathrm{nmean}_{tm} \, (\frac{\lambda_0}{\lambda_l} - 1)^m \ ) \tag{31}$$

$$\mathrm{ndiff}_{tl} = \sum_m (\ \mathrm{ndiff}_{tm} \, (\frac{\lambda_0}{\lambda_l} - 1)^m \ ) \tag{32}$$

where nmean_{tm} and ndiff_{tm} are the polynomial coefficients of order m of beam t, read from the DISP_MODEL.

From the dispersion and the metrology signal, we can compute the amount of delay introduced by the differential delay-line, for each spectral channel and each baseline. This quantity is called $\mathrm{OPD\_DISP}_{bfl}$:

$$\mathrm{OPD\_DISP}_{bfl} \ = \ \beta_{t_1\,l} \, \mathrm{OPD\_MET\_FC}_{t_1\,f} \ - \ \beta_{t_2\,l} \, \mathrm{OPD\_MET\_FC}_{t_2\,f} \ + \tag{33}$$

$$\gamma_{t_1\,l} \, \mathrm{FDDL}_{t_1\,f}/2 \ - \ \gamma_{t_2\,l} \, \mathrm{FDDL}_{t_2\,f}/2 \tag{34}$$

Dispersive group-delay and remaining phase

The signal $\mathrm{OPD\_DISP}_{bfl}$ contains a fraction of group-delay coded as a spectral slope. This shall be properly taken into account when attempting to compute astrometric/absolutes phases out of the dataset.

Therefore the pipeline also provide an additional representation of the same quantity, but decomposed into the total group-delay in the middle of the K-band ($\mathrm{GDELAY\_DISP}_{bf}$ in unit of distance, thus [m]), and the remaining phase ($\mathrm{PHASE\_DISP}_{lbf}$, in [rad]).

The total group-delay introduced by FDDL in the middle of the band is:

$$\mathrm{GDELAY\_DISP}_{bf} = \frac{\lambda_{l_1}^{-1} \, \mathrm{OPD\_DISP}_{l_1\,bf} \ - \ \lambda_{l_2}^{-1} \, \mathrm{OPD\_DISP}_{l_2\,bf}}{\lambda_{l_1}^{-1} - \lambda_{l_2}^{-1}} \tag{35}$$

where l_1 and l_2 are two consecutive wavelength channel in the middle of the band. The remaining phase is:

$$\mathrm{PHASE\_DISP}_{lbf} = \arctan(\ \exp(\frac{2i\pi}{\lambda_f} \, (\mathrm{OPD\_DISP}_{lbf} - \mathrm{GDELAY\_DISP}_{bf})) \ ) \tag{36}$$

The $\mathrm{GDELAY\_DISP}_{bf}$ quantities are defined to within a constant: the so called zeros Z_t (one per beam). When combined with the group-delays from SC and FT, the $\mathrm{GDELAY\_DISP}_{bf}$ can be used to construct a group-delay astrometry, e.g:

$$-2\pi \, \frac{\vec{B_{b_{ij}}}}{\lambda_l}.\vec{\delta} = \mathrm{GDELAY\_SC}_{b_{ij}f} - \mathrm{GDELAY\_FT}_{b_{ij}f} + \mathrm{GDELAY\_DISP}_{b_{ij}f} + \mathrm{Z_i} - \mathrm{Z_j} \tag{37}$$

A Installation

One is advised to read the installation instructions delivered with the GRAVITY pipeline distribution kit. These release-specific instructions can be found in the file `README` located in the top-level directory of the unpacked GRAVITY pipeline source tree. The supported platforms are listed in Section A.1. It is recommended reading through Section A.2.2 before starting the installation.

A bundled version of the GRAVITY pipeline with all the required tools and an installer script is available from `http://www.eso.org/pipelines/`, for users who are not familiar with the installation of software packages.

A.1 Supported platforms

The utilisation of the GNU build tools should allow to build and install the GRAVITY pipeline on a variety of UNIX platforms, but it has only been verified on the VLT target platforms:

- Linux (glibc 2.1 or later),
- Sun Solaris 2.8 or later,

using the GNU C compiler (version 3.2 or newer).

A.2 Building the GRAVITY pipeline

This section shows how to obtain, build and install the GRAVITY pipeline from the official source distribution.

A.2.1 Requirements

To compile and install the GRAVITY pipeline one needs:

- the GNU C compiler (version 3.2 or later),
- the GNU `gzip` data compression program,
- a version of the `tar` file-archiving program, and,
- the GNU `make` utility.

An installation of the Common Pipeline library (CPL) must also be available on the system. Currently the CPL version 2.1.1 or newer is required. The CPL distribution can be obtained from [3].

Please note that CPL itself depends on an existing qfits installation. The qfits sources are available from the CPL download page or directly from the qfits homepage at `http://www.eso.org/projects/aot/qfits`. In conjunction with CPL 2.1.1 qfits 5.3.1 must be used.

In order to run the GRAVITY pipeline recipes a front-end application is also required. Currently there are two such applications available, a command-line tool called *EsoRex* and the Java based data file organizer, *Gasgano*, which provides an intuitive graphical user interface (see Section 5.2, page 17). At least one of them must be installed. The *EsoRex* and *Gasgano* packages are available at `http://www.eso.org/cpl/esorex.html` and `http://www.eso.org/gasgano` respectively.

For installation instructions of any of the additional packages mentioned before please refer to the documentation of these packages.

A.2.2 Compiling and installing the GRAVITY pipeline

The GRAVITY pipeline distribution kit 1.0 contains:

gravity-manual-1.0.pdf	The GRAVITY pipeline manual
install_pipeline	Install script
cpl-7.0.tar.gz	CPL 7.0
esorex-3.12.3.tar.gz	esorex 3.12.3
gasgano-2.4.8.tar.gz	GASGANO 2.4.8for Linux
gravity-1.0.5.tar.gz	GRAVITY 1.0.5
gravity-calib-1.0.5.tar.gz	GRAVITY calibration files 1.0.5

Here is a description of the installation procedure:

1. Change directory to where you want to retrieve the GRAVITY pipeline recipes 1.0.5package. It can be any directory of your choice but not:

   ```
   $HOME/gasgano
   $HOME/.esorex
   ```

2. Download from the ESO ftp server, `http://www.eso.org/pipelines/`, the latest release of the GRAVITY pipeline distribution.

3. Verify the checksum value of the tar file with the cksum command.

4. Unpack using the following command:

   ```
   tar -xvf \pipename-kit-\pipelinevers.tar
   ```

   Note that the size of the installed software (including *Gasgano*) together with the static calibration data is about 1.3 Gb, mainly du to the test data.

5. Install: after moving to the top installation directory,

   ```
   cd \pipename-kit-\pipelinevers
   ```

it is possible to perform a simple installation using the available installer script (*recommended*):

```
./install_pipeline
```

(beware: the execution may take a few minutes on Linux and several minutes on SunOS).

Note that this release still needs to link to the eclipse library. At the end of the installation the user in addition to follow what reported by the installation script, needs to source an file ($HOME/..eclipse_bash.rc or $HOME/..eclipse_bash.rc, depending from the user shell) to set a few environment variables used by a few low level eclipse library based modules.

By default the script will install the GRAVITY recipes, *Gasgano*, *EsoRex*, all the necessary libraries, and the static calibration tables, into a directory tree rooted at `$HOME`. A different path may be specified as soon as the script is run.

The only exception to all this is the *Gasgano* tool, that will always be installed under the directory `$HOME/gasgano`. Note that the installer will move an existing `$HOME/gasgano` directory to `$HOME/gasgano.old` before the new *Gasgano* version is installed.

Important: the installation script would ensure that any existing *Gasgano* and *EsoRex* setup would be inherited into the newly installed configuration files (avoiding in this way any conflict with other installed instrument pipelines).

Alternatively, it is possible to perform a manual installation (*experienced users only*): the `README` file located in the top installation directory contains more detailed information about a step-by-step installation.

B Abbreviations and acronyms

ANSI	American National Standards Institute
ASCII	American Standard Code for Information Interchange
CalibDB	Calibration Database
CPL	Common Pipeline Library
DFO	Data Flow Operations department
DFS	Data Flow System department
DHS	Data Handling Server
DO	Data Organizer
DMD	Data Management and Operations Division
DRS	Data Reduction System
ESO	European Southern Observatory
ESOREX	ESO-Recipe Execution tool
FDDL	Fibered Differential Delay Lines
FITS	Flexible Image Transport System
FOV	Field Of View
FT	Fringe-Tracker
SC	Science-Combiner
GUI	Graphical User Interface
OB	Observation Block
OIFITS	OIFITS format, see [4]
PSO	Paranal Science Operations
QC	Quality Control
RON	Read Out Noise
SOF	Set Of Frames
UT	Unit Telescope
AT	Auxiliary Telescope
VLTI	Very Large Telescope Interferometer